College of Integrative Sciences and Arts
ARIZONA STATE UNIVERSITY

QCD THEORY SEMINAR
Dec 14, 2021
KEK

Chiral anomalous plasma in magnetospheres of pulsars

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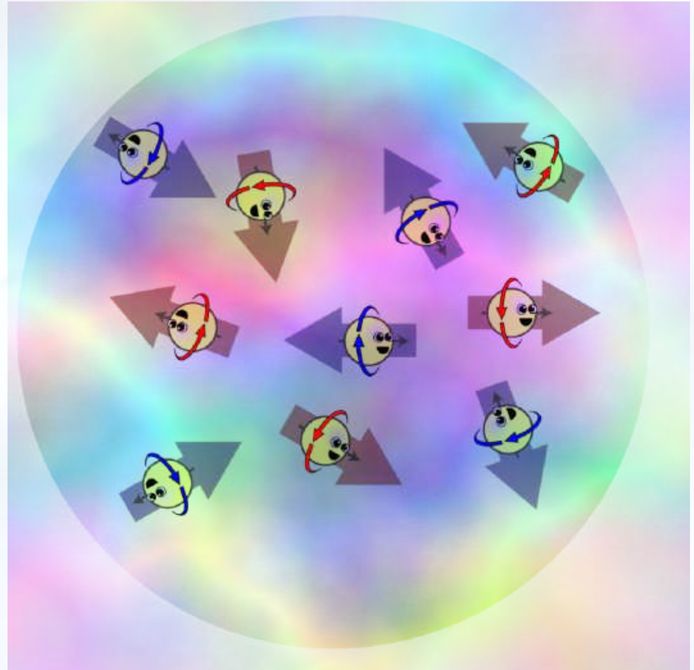
[E. V. Gorbar & I. A. Shovkovy, arXiv:2110.11380]

QCD theory seminar, Dec. 14, 2021



CHIRAL PLASMA

[Miransky & Shovkovy, Phys. Rep. 576, 1 (2015)]
[Kharzeev, Liao, Voloshin, Wang, Prog. Part. Nucl. Phys. 88, 1 (2016)]

Anomalous plasma

- Chiral relativistic plasma may allow $n_L \neq n_R$ to persist on *macroscopic* time/distance scales
- Slow evolution of $n = n_R + n_L$ and $n_5 = n_R - n_L$ is controlled by the continuity equations

$$\frac{\partial n}{\partial t} + \vec{\nabla} \cdot \vec{j} = 0$$

and

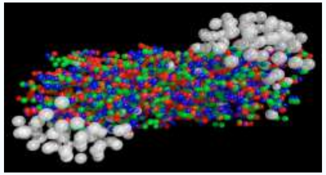
$$\frac{\partial n_5}{\partial t} + \vec{\nabla} \cdot \vec{j}_5 = \frac{e^2 \vec{E} \cdot \vec{B}}{2\pi^2} - \Gamma_{\rm m} n_5$$

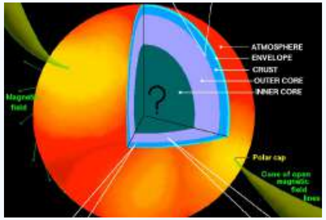
where the chirality flip rate: $\Gamma_{\rm m} \propto \alpha^2 T (m/T)^2$

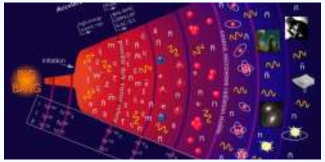
- Chiral anomaly can produce *macroscopic* effects in plasma

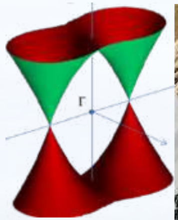
Chiral plasmas in nature

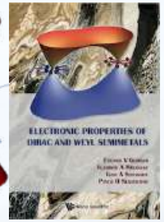
- **Heavy-ion collisions** **(high temperature)**
  [Kharzeev, Liao, Voloshin, Wang, Prog. Part. Nucl. Phys. 88, 1 (2016)]
- **Super-dense matter in compact stars** **(high density)**
  [Yamamoto, Phys. Rev. D **93**, 065017 (2016)]
- **Early Universe** **(high temperature)**
  [Boyarsky, Frohlich, Ruchayskiy, Phys. Rev. Lett. **108**, 031301 (2012)]
- **Electron plasma in Dirac/Weyl (semi-)metals**
  [Gorbar, Miransky, Shovkovy, Sukhachov, *Electronic Properties of Dirac and Weyl Semimetals* (World Scientific, Singapore, 2021)]
- **Other: cold atoms, superfluid ^{3}He-A, etc.**
  [Volovik, JETP Lett. 105, 34 (2017)]
- **Magnetospheres of magnetars** [Gorbar & Shovkovy, arXiv:2110.11380]
  (electron-positron plasma at moderately high temperature)











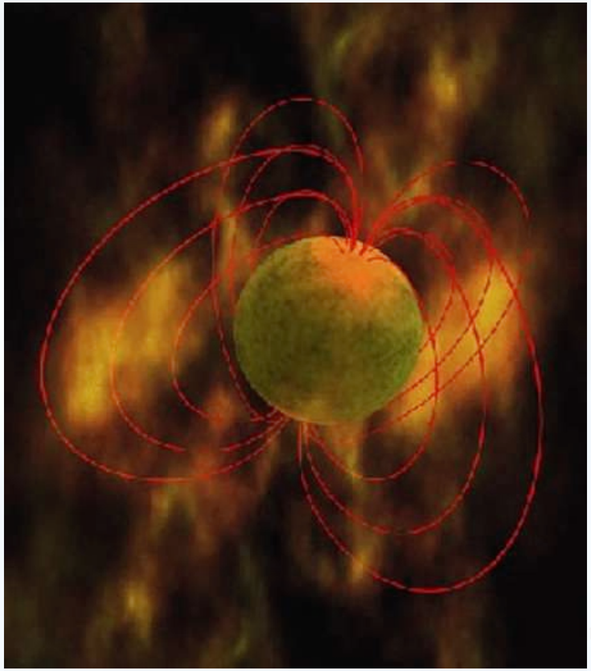


Image credit: NASA

PULSARS

Neutron stars

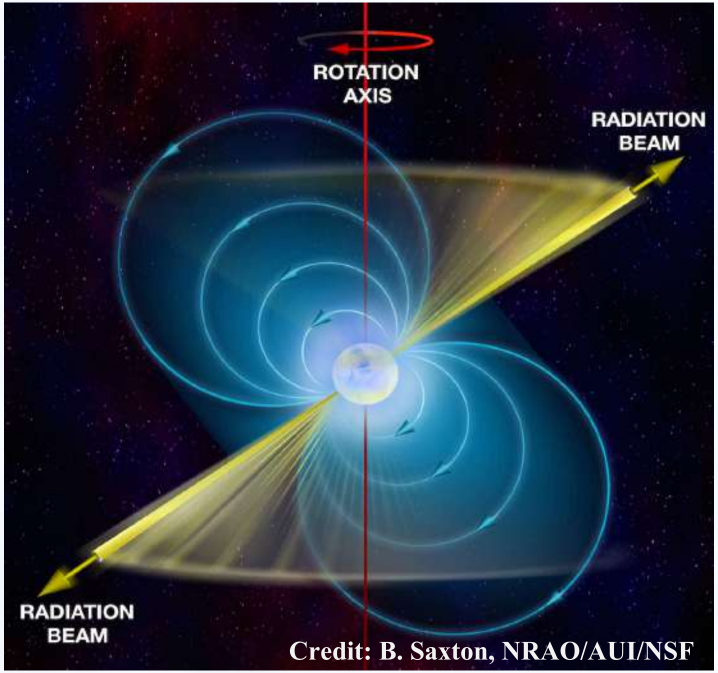
- **Neutron stars** are laboratories of matter under extreme conditions
- Prediction

[Baade & Zwicky, *Proc. Nat. Acad. Sci.* **20**, 259 (1934)]

- Observation

[Hewish, Bell, Pilkington, Scott & Collins, *Nature* **217**, 709 (1968)]

- **Pulsars** are neutron stars that are
  - rapidly rotating ($P \sim 1$ ms to 10 s)
  - strongly magnetized ($B \sim 10^8$ to 10^{15} G)



Credit: B. Saxton, NRAO/AUI/NSF

- Pulsar radiation is beamed along the magnetic field direction (the "lighthouse" effect)

Pulsars in P-$\dot{P}$ plane

- Characteristic age

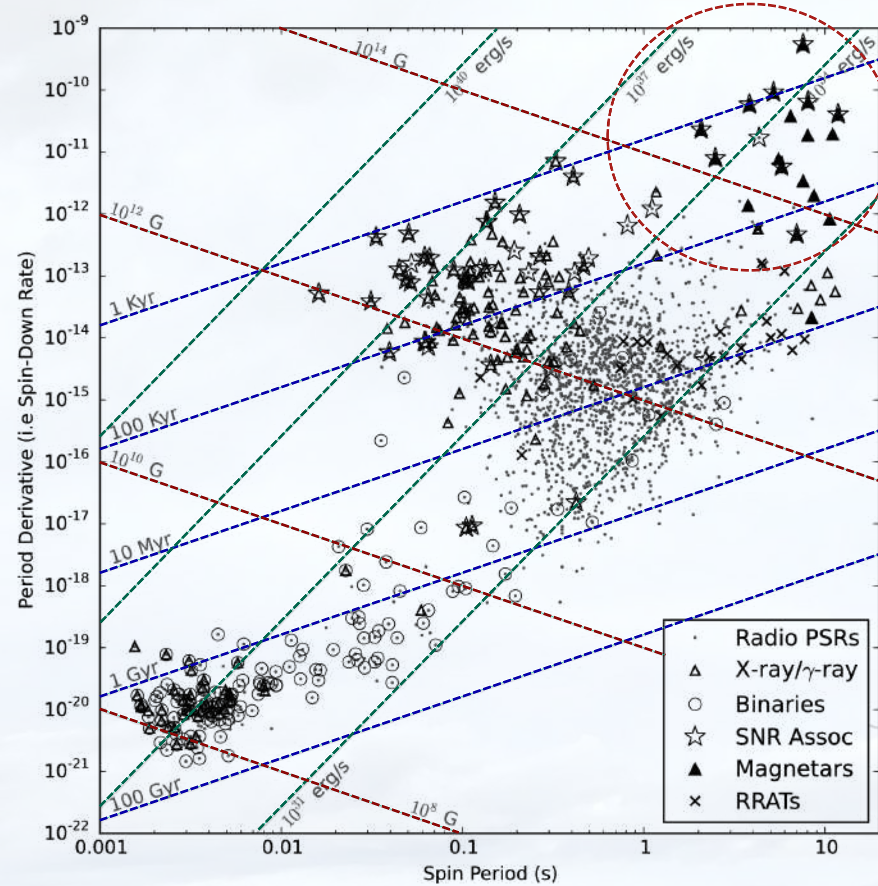
$$\tau \simeq \frac{P}{2\dot{P}}$$

- Spin-down luminosity

$$-\dot{E} \simeq 4\pi^2 I \frac{\dot{P}}{P^3}$$

- Characteristic magnetic field

$$B \simeq 3\times10^{19}\left(\frac{P\dot{P}}{s}\right)^{1/2}\mathrm{G}$$



J. Condon and S. Ransom, "Essential Radio Astronomy" (2016)

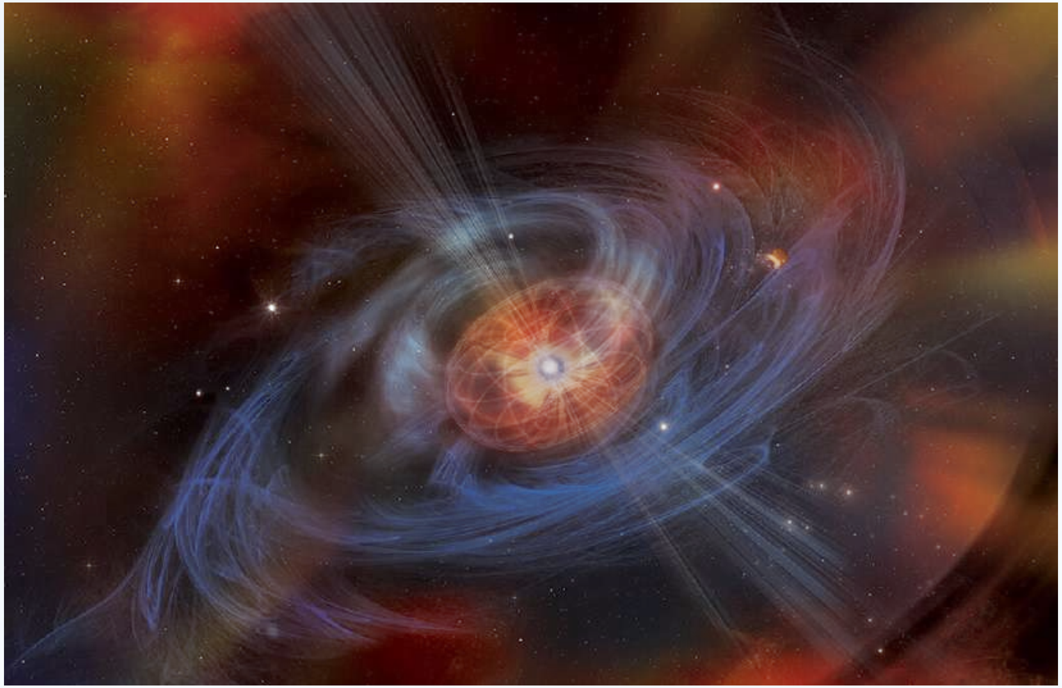


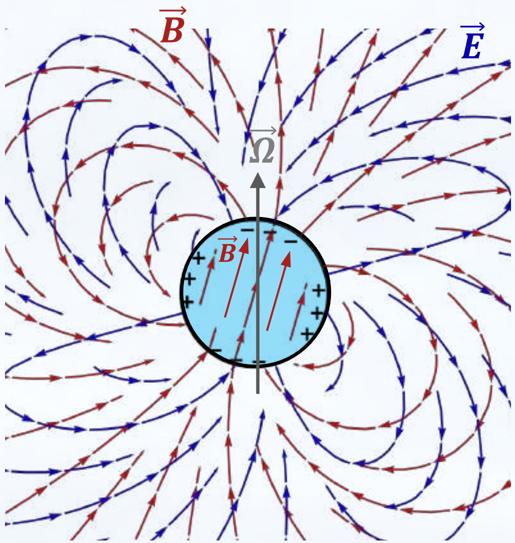
Image credit: Aurore Simonnet, Sonoma State University

MAGNETOSPHERES

Pulsar electrodynamics (VDM)

- **Vacuum dipole model (VDM)** ($\rho = 0$ & $J = 0$ outside the star)
- Stellar interior (good conductor):

$$\vec{E}'_{in} = \vec{E}_{in} + \frac{\vec{\Omega}\times\vec{r}}{c}\times\vec{B}_{in} = 0$$



- Fields outside the pulsar are

$$\vec{B} = \frac{B_0 R^3}{2r^3}(3(\hat{\boldsymbol{m}}\cdot\hat{\boldsymbol{r}})\hat{\boldsymbol{r}} - \hat{\boldsymbol{m}})$$

$$\vec{E} = \cdots$$ [see Deutsch, Ann. Astrophys. **18**, 1 (1955)]

where $\boldsymbol{m}$ is the magnetic moment and $\boldsymbol{\Omega}$ is the angular frequency

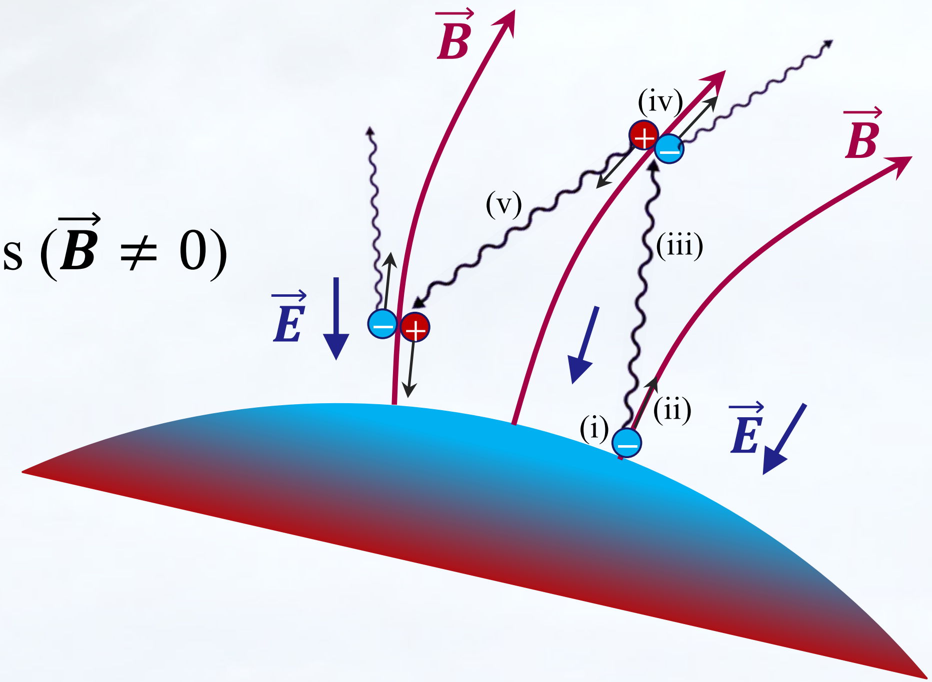
- There is a nonzero charge density and a strong electric field on the surface ($E_{\text{surf}} \sim \Omega R B_0 \sim 10^{12}$ to 10^{15} V/m)

Pulsar electrodynamics (VDM)

- Charged particles
  i. leave the surface ($\vec{E} \neq 0$)
  ii. move along curved trajectories ($\vec{B} \neq 0$)
  iii. produce curvature radiation
  iv. γ-quanta produce e^+e^- pairs

$$l_\gamma \simeq \frac{2R_c}{15}\frac{B_c}{B}\frac{m_e}{\varepsilon_\gamma}$$

  v. Secondary particles produce synchrotron & curvature radiation



- **End result:** (I) magnetized vacuum is nontransparent for photons with $\varepsilon_\gamma \gtrsim 2m_e$; (II) vacuum turns into plasma

Pulsar electrodynamics (RMM)

- **Rotating magnetosphere model (RMM)** (assuming highly conducting plasma outside the star)

$$\vec{\boldsymbol{E}}' = \vec{\boldsymbol{E}} + \frac{\vec{\boldsymbol{\Omega}}\times\vec{\boldsymbol{r}}}{c}\times\vec{\boldsymbol{B}} = 0$$

i.e., $E_{\parallel}=0$

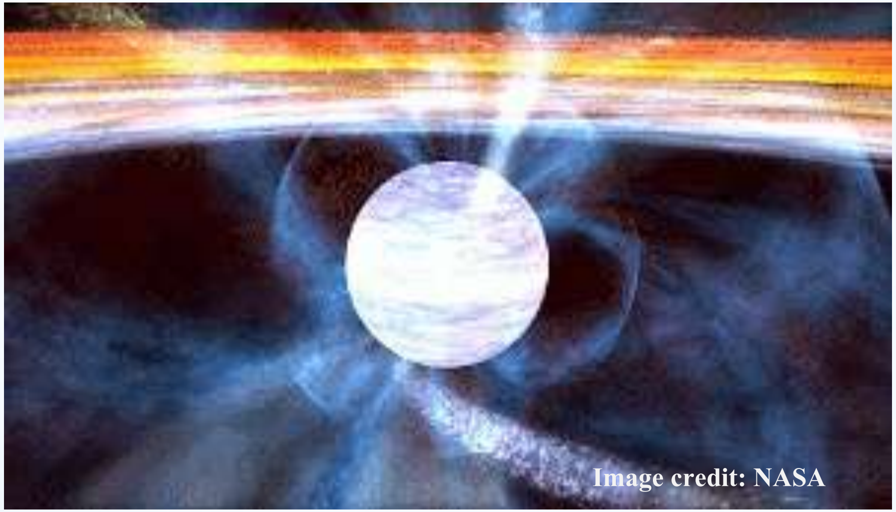
- Plasma motion is determined by

$$\vec{\boldsymbol{v}}_{\text{drift}} = c\frac{\vec{\boldsymbol{E}}\times\vec{\boldsymbol{B}}}{B^2} = \vec{\boldsymbol{\Omega}}\times\vec{\boldsymbol{r}} + j_{\parallel}\vec{\boldsymbol{B}}$$

- Corotating plasma is charged

$$\rho_{\text{GJ}} = \vec{\boldsymbol{\nabla}}\cdot\vec{\boldsymbol{E}} = -\frac{2}{c}\vec{\boldsymbol{\Omega}}\cdot\vec{\boldsymbol{B}}$$

[Goldreich & Julian, Astrophys. J. **157**, 869 (1969)]



Gaps in magnetosphere

- If one assumes that $E_{\|}=0$ everywhere, the magnetic field lines are equipotential ($V = \text{const}$)

- Then,

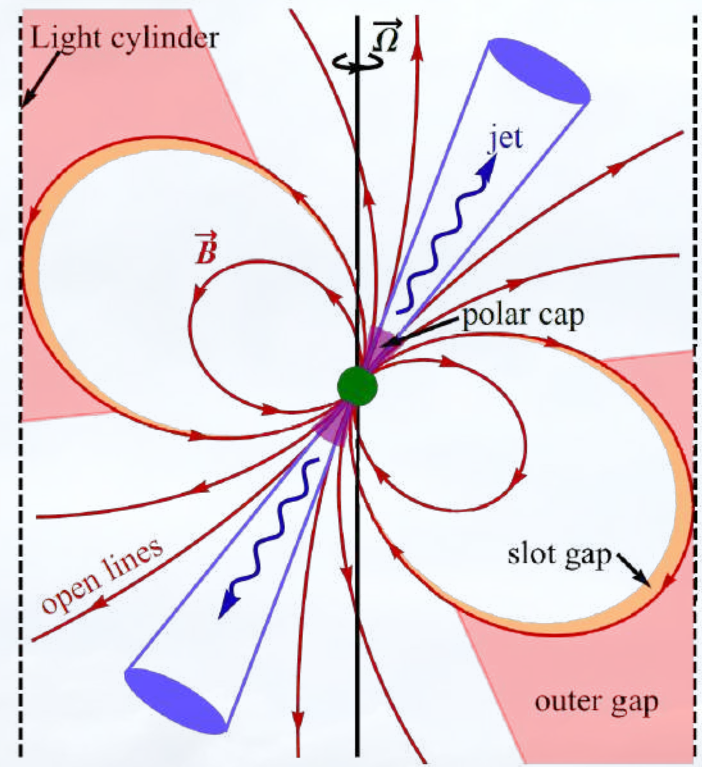
$$0 = \oint \vec{\boldsymbol{E}} \cdot d\vec{\boldsymbol{l}} = \int (\vec{\boldsymbol{\nabla}} \times \vec{\boldsymbol{E}}) \cdot d\vec{\boldsymbol{s}} = -\frac{\partial}{\partial t} \int \vec{\boldsymbol{B}} \cdot d\vec{\boldsymbol{s}}$$

- Thus, $E_{\|}=0$ cannot be enforced everywhere if $\vec{\boldsymbol{B}}$ changes in time
- Regions ("gaps") with unscreened $E_{\|}$ will necessarily develop (they result from dynamical charge/current starvation)

[Ruderman & Sutherland, Astrophys. J. **196**, 51 (1975)]

Gaps in magnetosphere

- Gaps can develop at various locations
- Intermittent gaps are caused by rapid outflow of charge
- The **gap size** h grows at a speed close to the speed of light
- Electric **potential** difference grows like $\Delta V = E_{\parallel} h \propto h^2$
- ΔV causes avalanche production of **electron-positron** pairs



[Ruderman & Sutherland, Astrophys. J. **196**, 51 (1975)]

- Since $B \propto 1/r^3$, anomalous effects are strongest near **polar caps**

Pulsar gaps

- Estimate of the gap size and the electric field

$$E_{\|} \simeq Bh/R_{LC}$$

where $R_{LC} = c/\Omega$ is the radius of light cylinder and

$$h \simeq 3.6 \text{ m} \left(\frac{R}{10 \text{ km}}\right)^{2/7} \left(\frac{\Omega}{1 \text{ s}^{-1}}\right)^{-3/7} \left(\frac{B}{10^{14} \text{ G}}\right)^{-4/7}$$

The field scales with pulsar parameters as follows

$$E_{\|} \approx 2.7\times10^{-8} E_c \left(\frac{R}{10 \text{ km}}\right)^{2/7} \left(\frac{\Omega}{1 \text{ s}^{-1}}\right)^{4/7} \left(\frac{B}{10^{14} \text{ G}}\right)^{3/7}$$

where $E_c = m_e^2/e = 1.3 \times 10^{18}$ V/m.

[Ruderman & Sutherland, Astrophys. J. **196**, 51 (1975)]

Gap parameters

- Quantitative estimate of the gap size and fields

B	10^{12} G	10^{13} G	10^{14} G	10^{15} G
h	50 m	13.4 m	3.6 m	0.97 m
$\frac{E_\parallel}{E_c}$	3.8×10^{-9}	1.0×10^{-8}	2.7×10^{-8}	7.3×10^{-8}
$\frac{\boldsymbol{E}\cdot\boldsymbol{B}}{E_c B_c}$	8.6×10^{-11}	2.3×10^{-9}	6.2×10^{-8}	1.7×10^{-6}

where

$$E_c = m_e^2/e = 1.3 \times 10^{18}\ \text{V/m}$$

$$B_c = m_e^2/e = 4.4 \times 10^{13}\ \text{G}$$

Chiral charge production

- The evolution of the chiral charge is determined by

$$\frac{\partial n_5}{\partial t} + \vec{\nabla}\cdot\vec{j}_5 = \frac{e^2\vec{E}\cdot\vec{B}}{2\pi^2} - \Gamma_{\rm m} n_5$$

- While the chiral anomaly produces n_5, the chirality flipping tries to wash it away
- The chiral charge n_5 approaches the following steady-state value:

$$n_5 = \frac{e^2}{2\pi^2\Gamma_{\rm m}}\vec{E}\cdot\vec{B}$$

- The estimates for the chirality flip rate in a hot plasma

$$\Gamma_{\rm m} \simeq \frac{\alpha^2 m_e^2}{T} \quad (T \lesssim m_e/\sqrt{\alpha}) \quad \text{and} \quad \Gamma_{\rm m} \simeq \frac{\alpha\, m_e^2}{T} \quad (T \gg m_e/\sqrt{\alpha})$$

[Boyarsky, Cheianov, Ruchayskiy, Sobol, Phys. Rev. Lett. **126**, 021801 (2021)]

Time scales

- The gap formation time

$$t_h \sim h/c \sim 10^{-8}\ \mathrm{s}$$

- Timescale for chiral charge production

$$t^{\star} \sim 1/\Gamma_{\mathrm{m}} \sim 10^{-17}\ \mathrm{s}$$

- Note that

$$t_h \gg t^{\star}$$

- Thus, the chirality production is nearly instantaneous

Estimate for n_5 in magnetars

- The estimate for the chiral charge is given by

$$n_5 \simeq \frac{e^2 E_{\|} B}{2\pi^2 \Gamma_m} \simeq 1.5 \times 10^{-5}\ \mathrm{MeV}^3 \left(\frac{T}{1\ \mathrm{MeV}}\right) \times \left(\frac{R}{10\ \mathrm{km}}\right)^{2/7} \left(\frac{\Omega}{1\ \mathrm{s}^{-1}}\right)^{4/7} \left(\frac{B}{10^{14}\ \mathrm{G}}\right)^{10/7}$$

- The corresponding chiral chemical potential is

$$\mu_5 \simeq \frac{3 n_5}{T^2} \simeq 4.6 \times 10^{-5}\ \mathrm{MeV} \left(\frac{T}{1\ \mathrm{MeV}}\right)^{-1} \times \left(\frac{R}{10\ \mathrm{km}}\right)^{2/7} \left(\frac{\Omega}{1\ \mathrm{s}^{-1}}\right)^{4/7} \left(\frac{B}{10^{14}\ \mathrm{G}}\right)^{10/7}$$

Values of n_5 and μ_5

- The corresponding numerical values for chiral charge and chiral chemical potential are

B	10^{12} G	10^{13} G	10^{14} G	10^{15} G
h	50 m	13.4 m	3.6 m	0.97 m
$\frac{E_{\parallel}}{E_c}$	3.8×10^{-9}	1.0×10^{-8}	2.7×10^{-8}	7.3×10^{-8}
$\frac{\boldsymbol{E}\cdot\boldsymbol{B}}{E_c B_c}$	8.6×10^{-11}	2.3×10^{-9}	6.2×10^{-8}	1.7×10^{-6}
$\frac{n_5}{m_e^3}$	1.6×10^{-7}	4.3×10^{-6}	1.1×10^{-4}	3.1×10^{-3}
$\frac{\mu_5}{m_e}$	1.2×10^{-7}	3.4×10^{-6}	9.0×10^{-5}	2.4×10^{-3}

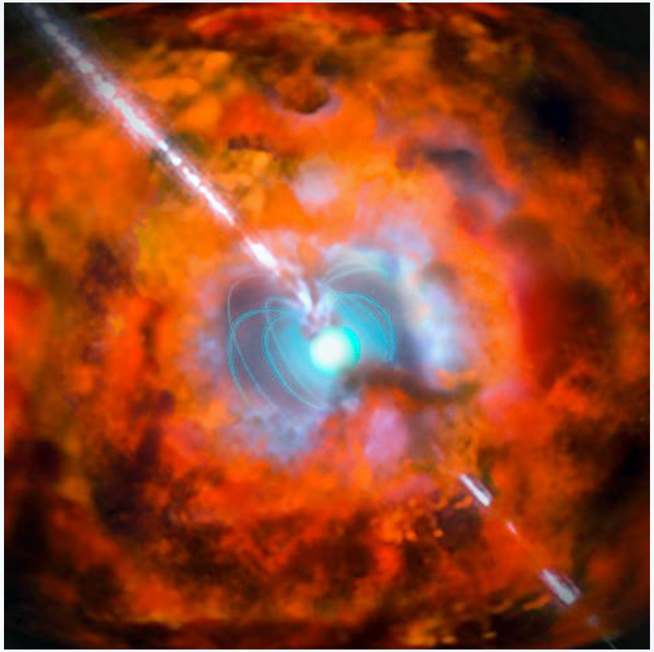


Image credit: European Southern Observatory

CHIRAL PLASMA INSTABILITY

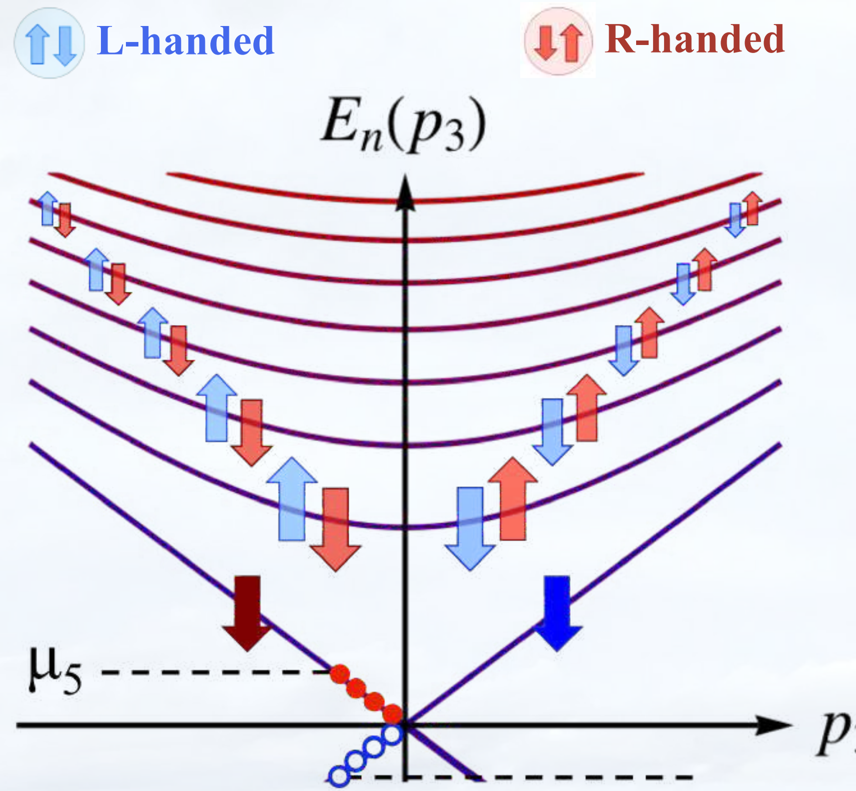
[]

Plasma with $\mu_5 \neq 0$

- Nonzero μ_5 and $\vec{B}$ drive the chiral magnetic effect

$$\vec{j} = \frac{e^2 \vec{B}}{2\pi^2} \mu_5$$

- The effect comes from the spin-polarized LLL (s=↓)
  - L-handed states ($p_3 < 0$ & $|E| < \mu_5$) are empty (holes with $p_3 > 0$)
  - R-handed states ($p_3 < 0$ & $E < \mu_5$) are occupied



- However, plasma at $\mu_5 \neq 0$ is unstable

Maxwell equations at $\mu_5 \neq 0$

- The total current (CME + Ohm)

$$\boldsymbol{j} = \frac{2\alpha}{\pi}\mu_5 \mathbf{B} + \sigma \mathbf{E}$$

- By substituting $\boldsymbol{j}$ into Ampere's law

$$\boldsymbol{\nabla} \times \mathbf{B} = \boldsymbol{j} + \frac{\partial \mathbf{E}}{\partial t}$$

and solving for the electric field, one derives

$$\mathbf{E} = \frac{1}{\sigma}\left(\boldsymbol{\nabla} \times \mathbf{B} - k_\star \mathbf{B} - \frac{\partial \mathbf{E}}{\partial t}\right)$$

where $k_\star = \frac{2\alpha\mu_5}{\pi}$

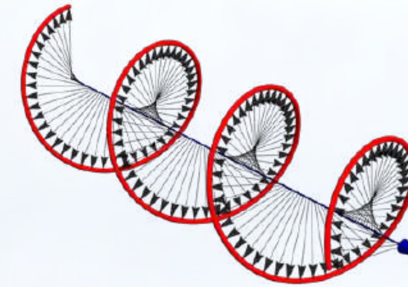
- Finally, by using Faraday's law, one has

$$\frac{\partial \mathbf{B}}{\partial t} = -\frac{1}{\sigma}\left(\boldsymbol{\nabla} \times (\boldsymbol{\nabla} \times \mathbf{B}) - k_\star \boldsymbol{\nabla} \times \mathbf{B} + \frac{\partial^2 \mathbf{B}}{\partial t^2}\right)$$

Helical modes $\mu_5 \neq 0$

- Search for a solution as a superposition of helical eigenstates

$$\boldsymbol{\nabla} \times \mathbf{B}_{\lambda,k} = \lambda k \mathbf{B}_{\lambda,k}$$



e.g.,

$$\mathbf{B}_{\lambda,k} = B_0 \left(\hat{\boldsymbol{x}} + i\lambda \hat{\boldsymbol{y}}\right) e^{-i\omega t + ikz}$$

Then, for a fixed eigenmode, the evolution equation reads

$$\frac{d\mathbf{B}_{\lambda,k}}{dt} = \frac{1}{\sigma}\left(\lambda k_\star k - k^2 - \frac{\partial^2}{\partial t^2}\right)\mathbf{B}_{\lambda,k}$$

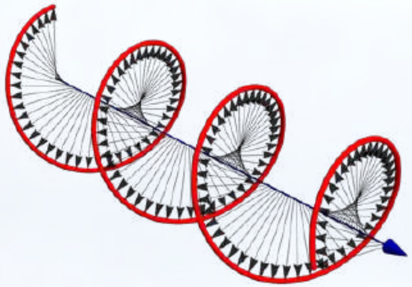
- The two solutions for the frequency are

$$\omega_{1,2} \;=\; -\frac{i}{2}\left(\sigma \pm \sqrt{\sigma^2 + 4k(\lambda k_\star - k)}\right)$$

Long-wavelength modes

- For plasma with high conductivity

$$\omega_{1,2} \simeq \begin{cases} -i\left(\sigma + \frac{k(\lambda k_\star - k)}{\sigma}\right) \\ i\frac{k(\lambda k_\star - k)}{\sigma} \end{cases}$$



- The 1st mode is damped by charge screening:

$$B_{k,1} \propto B_0 e^{-\sigma t}$$

- The 2nd mode is unstable when $k < \lambda k_\star$:

$$B_{k,2} \propto B_0 e^{+tk(\lambda k_\star - k)/\sigma}$$

- The momentum of the fastest growing mode $B_{k,2}$ is

$$\frac{1}{2} k_\star$$

Instability in pulsars

- The estimate for $k_\star$

$$k_\star \simeq 2.2\times10^{-7}\ \mathrm{MeV}\left(\frac{T}{1\ \mathrm{MeV}}\right)^{-1}\left(\frac{R}{10\ \mathrm{km}}\right)^{2/7}\left(\frac{\Omega}{1\ \mathrm{s}^{-1}}\right)^{4/7}\left(\frac{\mathrm{B}}{10^{14}\ \mathrm{G}}\right)^{10/7}$$

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$\frac{n_5}{m_e^3}$	1.6×10^{-7}	4.3×10^{-6}	1.1×10^{-4}	3.1×10^{-3}
$\frac{\mu_5}{m_e}$	1.2×10^{-7}	3.4×10^{-6}	9.0×10^{-5}	2.4×10^{-3}
$\frac{k_\star}{m_e}$	5.8×10^{-10}	1.6×10^{-8}	4.2×10^{-7}	1.1×10^{-5}

Observational consequences

- Unstable plasma in the gaps produces **helical modes** in the frequency range

$$0 \lesssim \omega \lesssim k_\star$$

- For magnetars, these span **radio frequencies** and may reach into the **near-infrared** range
- Key characteristics: **circularly** polarized radio emission
- Available energy

$$E_{\text{tot}} \sim \mu_5^2 T^2 h^3 \sim 10^{23}\ \text{erg} \quad \text{to} \quad 10^{28}\ \text{erg}$$

- The energy is sufficient to feed the **fast radio bursts** (FRB)

Outstanding problems

- Interplay of chiral charge and electron-positron pair **production** induced by energetic photons should be studied in detail
- The modification of the **chiral flip rate** $\Gamma_{\mathrm{m}} \simeq \frac{\alpha^2 m_e^2}{T}$ by the strong magnetic field (extra suppression?)
- The role of the **inverse magnetic cascade** and the **chiral-magnetic turbulence** should be quantified
- Self-consistent **dynamics** of chiral plasma in the gap regions should be simulated in detail
- Detailed mechanism of the **energy transfer** from unstable helical modes to radio emission in FRBs

Summary

- Chiral anomaly can have *macroscopic* implications in pulsars
- It leads to a *significant* chiral charge production (up to 10^{34} m^{-3}) in strongly magnetized magnetospheres
- The chiral chemical potential μ_5 can be up to 10^{-3} MeV
- This is sufficient to trigger emission of helical waves with frequencies up to about $k_\star \simeq \frac{2}{\pi}\alpha\mu_5$ (radio to infrared range)
- Helical waves can affect the pulsar jets and observable features of the fast radio bursts
- For quantitative effects, further detailed studies are needed